

Cosmology with a light ghost

Mikhail M. Ivanov^{a,b,c} and Anna A. Tokareva^{a,b}

^aInstitute of Physics, LPPC, École Polytechnique Fédérale de Lausanne,
CH-1015, Lausanne, Switzerland

^bInstitute for Nuclear Research of the Russian Academy of Sciences,
60th October Anniversary Prospect, 7a, 117312 Moscow, Russia

^cDepartment of Particle Physics and Cosmology, Physics Faculty, Moscow State University,
Vorobjevy Gory, 119991, Moscow, Russia

E-mail: mikhail.ivanov@cern.ch, anna.tokareva@epfl.ch

Received October 19, 2016

Revised November 28, 2016

Accepted December 2, 2016

Published December 12, 2016

Abstract. We study the creation and evolution of cosmological perturbations in renormalizable quadratic gravity with a Weyl term. We adopt a prescription that implies the stability of the vacuum at the price of introducing a massive spin-two ghost state, leading to the loss of unitarity. The theory may still be predictive regardless the interpretation of non-unitary processes provided that their rate is negligible compared to the Universe expansion rate. This implies that the ghost is effectively stable. In such a setup, there are two scalar degrees of freedom excited during inflation. The first one is the usual curvature perturbation whose power spectrum appears to coincide with that of single-field inflation. The second one is a scalar component of the ghost encoded in the shift vector of the metric in the uniform inflaton gauge. The amplitudes of primordial tensor and vector perturbations are strongly suppressed. After inflation the ghost field starts to oscillate and its energy density shortly becomes dominant in the Universe. For all ghost masses allowed by laboratory constraints ghosts should have “overclosed” the Universe at temperatures higher than that of primordial nucleosynthesis. Thus, the model with the light Weyl ghost is ruled out.

Keywords: inflation, modified gravity, physics of the early universe, quantum gravity phenomenology

ArXiv ePrint: [1610.05330](https://arxiv.org/abs/1610.05330)



Contents

1	Introduction	1
2	Preliminaries	3
3	Inflation	5
3.1	Exact de Sitter background	6
3.2	Beyond the limit of exact de Sitter space	10
3.3	Observational predictions	11
4	After inflation: the cosmological ghost problem	12
5	Discussion and conclusions	15
A	Gauge transformations	16
B	Inflationary perturbations in the uniform inflaton gauge	16
C	Cosmological perturbations at late times	18
C.1	Main equations: qualitative analysis	18
C.2	Ghost energy density	20

1 Introduction

General Relativity (GR) is a very powerful effective field theory whose predictions were verified at various scales. However, elegant as it is, GR cannot be a fundamental theory of gravity since it is not able to account for physics above the Planck energy. The quest for a consistent theory of quantum gravity is not yet completed, although there exist promising approaches to this problem. String theory (e.g. [1]) provides important insights at the price of great complexity, which makes challenging the investigation of observational consequences.

An interesting proposal is Lorentz-violating Hořava-Lifshitz gravity [2], a branch of which (the so-called projectable model) provides the only known example of a consistent quantum field theory for the spin-2 graviton in four dimensions [3]. Unfortunately, this branch does not reproduce GR at low energies within weak coupling [4, 5]. Another, the so-called non-projectable branch of Hořava-Lifshitz gravity [6] is not yet rigorously proven to be renormalizable but is consistent with all current observations [7, 8] and has curious implications for inflation [9], dark matter [10], and dark energy [11], see [12] for a review.

Another model of quantum gravity has been known for decades. This is the gravity theory with quadratic curvature invariants developed by Stelle in refs. [13, 14]. For some regions of parameter space quadratic gravity even becomes asymptotically free and thus UV complete [15–17]. Because of the presence of negative-norm states in the spectrum this theory was not very popular, but regained interest recently [18–20].

Quadratic gravity is a higher-derivative theory, which is known to be plagued by the Ostrogradsky instability at the classical level [21, 22]. At the quantum level the theory possesses a spin-2 state with a negative kinetic term. The usual quantization prescription implies that the energy of this state is unbounded from below, which triggers an inevitable vacuum

decay. However, the spin-2 state can be quantized in a way that avoids this instability by introducing a negative norm [23]; the resulting particle is called “ghost” and has a positively-defined energy. The presence of the ghost does not allow the theory to be interpreted along the lines of unitary quantum mechanics, however, a consistent interpretation might still exist, see e.g. refs. [24] and references therein. First, ghosts may be removed from the partition function by some non-trivial boundary conditions [25]. Second, negative-norm states may be not so dangerous if for all measurable outcomes of an experiment overall probabilities remain positive and not exceeding unity [26]. On the other hand, one can think of quadratic gravity as an “intermediate” effective theory with a hope that it admits a unitary UV-completion.

In this study we adopt an agnostic point of view about this problem: we do not know the interpretation of non-unitary processes, but quadratic gravity still can be predictive regardless any interpretation provided that these processes are negligible during the lifetime of the Universe. In this case the violation of unitarity, even if observable in principle, is suppressed, and thus the theory becomes viable, at least, at the phenomenological level.

As shown below, this setup requires the ghost mass to be smaller than 10 MeV. This implies that the ghost has non-negligible fluctuations at inflation, which might have observational consequences.

Inflationary perturbations in quadratic gravity in the case of the ghost mass bigger than the Hubble rate were studied in ref. [27]. Tensor modes were discussed in detail in ref. [28] for a generic ghost mass. This analysis showed that the amplitude of tensor modes is negligible in the limit of a vanishing ghost mass. Inflation in Weyl gravity has been analyzed in [29, 30], where it was also shown that amplitudes of tensor and vector perturbations are largely suppressed in the limit of a small ghost mass. In ref. [29] it was also found that mode functions of the Newtonian gravitational potential are unstable at superhorizon scales, but no interpretation of this fact was given.

In this paper we investigate the dynamics of cosmological perturbations in renormalizable quadratic gravity both at inflation and late times. We show that in the conformal Newtonian gauge the ghost field acquires a tachyonic mass on the de Sitter background, which leads to a superhorizon growth of ghost fluctuations during inflation. The back-reaction of metric perturbations is not negligible and it induces a superhorizon growth of inflaton perturbations. We argue that this instability is a gauge artifact since there is at least one gauge (namely, the uniform inflaton gauge) free of superhorizon growth.

We find two scalar modes that get excited and then conserve outside the horizon during inflation. The first one is the scalar curvature perturbation, whose amplitude turns out to coincide with the standard single-field expression. The second one is a longitudinal part of the shift vector, which is related to the ghost degree of freedom. At late times, once the Hubble rate gets equal to the ghost mass, this mode starts to oscillate and the corresponding energy density shortly becomes dominant in the Universe. The situation is somewhat close to the cosmological “light moduli problem” [31], however, in our case there are crucial differences that make the problem more severe. Our main conclusion is that cosmology with a light ghost is incompatible with a healthy late-time evolution.

The paper is organized as follows. In section 2 we discuss quadratic gravity and the constraints on the ghost mass. In section 3 we explore the generation of scalar perturbations at inflation and observational signatures of the model. In the section 4 we study the late-time evolution and elaborate on the ghost domination. We discuss our results and conclude in section 5. In appendix A we introduce the relations between perturbations in the conformal Newtonian and the uniform inflaton gauges. In appendix B we cross-check the results of the

main text in the uniform inflaton gauge. Finally, in appendix C we study qualitatively the equations for cosmological perturbations at late times, and derive the ghost energy density.

2 Preliminaries

We start with the action of quadratic renormalizable gravity,

$$S = \frac{M_P^2}{2} \int d^4x \sqrt{-g} \left(R + \frac{1}{6\mu^2} R^2 - \frac{1}{2m^2} C_{\mu\nu\rho\sigma} C^{\mu\nu\rho\sigma} \right), \quad (2.1)$$

where R is the scalar curvature, $C_{\mu\nu\rho\sigma}$ is the Weyl tensor and M_P is the reduced Planck mass.¹ Upon flat space this theory possesses a massless transverse-traceless spin-2 excitation (graviton), a spin-0 particle of mass μ (dubbed scalaron [32, 33]) and a spin-2 particle of mass m with the “wrong” sign of its quadratic lagrangian. In what follows, we adopt a negative-norm prescription for this state, which implies the loss of unitarity but ensures the stability with respect to the vacuum decay. In this paper we assume that the ghost and scalaron masses are not tachyonic,

$$m^2 > 0 \quad \text{and} \quad \mu^2 > 0, \quad (2.2)$$

i.e. both degrees of freedom have standard dispersion relations upon flat space,²

$$\begin{aligned} \text{scalaron:} \quad \omega^2 &= k^2 + \mu^2, \\ \text{ghost:} \quad \omega^2 &= k^2 + m^2. \end{aligned} \quad (2.3)$$

Although the condition $\mu^2 > 0$ corresponds to a region of parameter space incompatible with asymptotic freedom [18, 19], the case of $\mu^2 < 0$ is not viable as the cosmological evolution leads either to a collapse or to a run-away singularity [34–36].

It is convenient to switch to the Einstein frame by means of a conformal transformation [37]. Recall that the square of the Weyl tensor is an invariant under the conformal symmetry. This yields the following action:

$$\begin{aligned} S &= S_{GR} + S_\phi + S_{\text{Weyl}} \\ &= \frac{M_P^2}{2} \int d^4x \sqrt{-g} R - \frac{1}{2} \int d^4x \sqrt{-g} (\partial_\mu \phi \partial^\mu \phi + 2V(\phi)) - \frac{M_P^2}{4m^2} \int d^4x \sqrt{-g} C_{\mu\nu\rho\sigma} C^{\mu\nu\rho\sigma}, \end{aligned} \quad (2.4)$$

where $V(\phi)$ is the Starobinsky potential,

$$V(\phi) = \frac{3}{4} \mu^2 M_P^2 \left(1 - e^{-\sqrt{\frac{2}{3}} \frac{\phi}{M_P}} \right)^2. \quad (2.5)$$

Although the Starobinsky model [32, 33]³ provides us with the most natural way to organize inflation in quadratic gravity, in our analysis we will assume a generic potential $V(\phi)$ with the requirements that it allows for a slow-roll regime and may be obtained from a renormalizable

¹Units: $M_P = 1/\sqrt{8\pi G} \simeq 2.4 \times 10^{18}$ GeV, $c = 1$. Conventions: $(-+++)$; $R^\mu{}_{\nu\rho\sigma} = \partial_\rho \Gamma^\mu_{\nu\sigma} - \partial_\sigma \Gamma^\mu_{\nu\rho} + \dots$; $R_{\nu\sigma} = R^\mu{}_{\nu\mu\sigma}$; $R = g^{\mu\nu} R_{\mu\nu}$; $G_{\mu\nu} = R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R$. $C_{\mu\nu\rho\sigma} = R_{\mu\nu\rho\sigma} - \frac{1}{2} (g_{\mu\rho} G_{\nu\sigma} - g_{\mu\sigma} G_{\nu\rho} - g_{\nu\rho} G_{\mu\sigma} + g_{\nu\sigma} G_{\mu\rho}) - \frac{R}{3} (g_{\mu\rho} g_{\nu\sigma} - g_{\mu\sigma} g_{\nu\rho})$. Greek indices run from 0 to 3; latin indices run from 1 to 3.

²This means that the mass term of the ghost should have the sign opposite to that of a healthy bosonic field.

³For the dynamics in the Einstein frame see ref. [38].

theory. In this way our results will be more model-independent, and applicable, for instance, to agravity-based inflationary models discussed in refs. [18, 39]. Note that if we had not required renormalizability, the model with a generic potential would, formally, be equivalent to $f(R)$ -gravity plus a Weyl term. We will refer to field ϕ as “inflaton” in what follows.

The variation of (2.4) with respect to the metric and ϕ yields the system of the modified Einstein and Klein-Gordon-Fock equations,

$$\begin{aligned} \nabla^\mu \nabla_\mu \phi - V'(\phi) &= 0, \\ R_{\mu\nu} - \frac{1}{2} R g_{\mu\nu} - \frac{1}{m^2} B_{\mu\nu} &= \frac{1}{M_P^2} T_{\mu\nu}, \\ \text{where } B_{\mu\nu} &\equiv \left(2\nabla^\alpha \nabla^\beta + R^{\alpha\beta} - \frac{1}{2} g^{\alpha\beta} R \right) C_{\mu\alpha\nu\beta}, \\ T_{\mu\nu} &= \partial_\mu \phi \partial_\nu \phi - g_{\mu\nu} \left(\frac{1}{2} (\partial_\mu \phi)^2 + V(\phi) \right), \end{aligned} \tag{2.6}$$

and ∇_μ is the covariant derivative.

The presence of negative-norm states implies processes with negative widths and cross-sections. Their interpretation is still an open issue, e.g. a negative decay width might actually signalize the violation of causality [40]. As shown in this reference, even unitarity may be restored within the acausal interpretation. It is not clear, however, how to extend this interpretation to account for non-unitary processes in cosmology. On the other hand, one can try to adopt the “usual” probabilistic interpretation, which suggests that negative-probability processes are instabilities. Indeed, in such a setup the “ghost decay” implies an exponential growth of ghost particles’ energy density, which follows from the Boltzmann equation,

$$\frac{dn_{gh}}{dt} + 3Hn_{gh} = |\Gamma|n_{gh}, \tag{2.7}$$

where H is the Hubble rate, n_{gh} is the ghost number density, and Γ is the (negative) decay rate of ghosts.⁴

The scalaron and the ghost are identified with metric components and thus couple to matter universally (see e.g. ref. [41]). The expression for the decay width of ghosts is similar (up to a numerical factor) to that of the scalaron [42, 43], but with a minus sign originating from the negative-norm prescription,

$$\Gamma \sim -\frac{m^3}{M_P^2}. \tag{2.8}$$

One can suppress this process by properly tuning the ghost mass. Requiring the characteristic time of ghost decay to be larger than the lifetime of the Universe one gets a conservative bound

$$m \lesssim 10 \text{ MeV}. \tag{2.9}$$

The lower bound on the ghost mass comes from short-scale gravity experiments. In the Newtonian limit the action (2.1) yields the usual GR gravitational potential plus corrections

⁴It should be pointed out that there can be many other ways to interpret theories with negative-norm states [24]. For instance, one can define probabilities to be always positive, but then their sum will exceed unity. In this case negative states do not imply negative probabilities.

suppressed by Yukawa-type exponents [13]. The bounds on the latter [44, 45] imply the following double-sided constraint on the ghost mass:

$$10^{-1} \text{ eV} \lesssim m \lesssim 10 \text{ MeV}. \quad (2.10)$$

To summarize, let us outline different regimes of quadratic gravity depending on the ghost mass m and its relation to the Hubble rate at inflation H_{inf} ,

- $m > H_{\text{inf}}$ (“heavy ghost”). The ghost is gapped at inflation and has no effect on cosmology. Note that adimensional gravity requires the ghost mass to be smaller than $3 \times 10^{10} \text{ GeV}$ for the naturalness of the small electroweak scale [18]. Since for simplest inflationary scenarios $H_{\text{inf}} \gtrsim 10^{13} \text{ GeV}$, achieving the naturalness might be challenging in agravity with a “heavy ghost”.⁵
- $10 \text{ MeV} < m \ll H_{\text{inf}}$ (“intermediate mass ghost”). The ghost is gapless at inflation and thus its fluctuations get excited. For this mass range non-unitary processes (like the ghost decay) happen after inflation and thus the viability of the theory depends on their interpretation. Since a consistent interpretation does not exist at the moment, phenomenological implications are speculative. Recall that if $m \lesssim 10^{10} \text{ GeV}$, the small electroweak scale is natural in agravity. Note that our analysis of inflation (section 3) is applicable to this mass range.
- $0.1 \text{ eV} \lesssim m \lesssim 10 \text{ MeV}$ (“light ghost”). The ghost decay is suppressed and the laboratory constraints are satisfied. Quadratic gravity may be phenomenologically viable and its predictions are not sensitive to any interpretation of non-unitary processes. We will focus on this regime and study the theory both at inflation (section 3) and late times (section 4).

3 Inflation

We start with the creation of cosmological perturbations at inflation. Since the Weyl term disappears on the homogeneous and isotropic metric, the background inflationary dynamics in our model is similar to that of single-field inflation. By plugging the homogeneous and isotropic ansatz,

$$\begin{aligned} ds^2 &= a(\eta)^2(-d\eta^2 + \mathbf{d}\mathbf{x}^2), \\ \phi &= \phi(\eta), \end{aligned} \quad (3.1)$$

into eq. (2.6), we obtain the usual background equations:

$$\begin{aligned} 3M_P^2 \mathcal{H}^2 &= \frac{\phi'^2}{2} + a^2(\eta)V, \\ 3M_P^2 \mathcal{H}' &= -\phi'^2 + a^2(\eta)V, \\ \phi'' + 2\mathcal{H}\phi' + a^2 V_{,\phi} &= 0, \end{aligned} \quad (3.2)$$

where primes denote partial derivatives w.r.t. conformal time, $' \equiv \partial/\partial\eta$, and $\mathcal{H} \equiv a'/a$ is the conformal Hubble parameter. Where convenient, we will also use the quantities expressed

⁵Among all of scenarios studied in ref. [39], the lowest possible Hubble scale at inflation is given by that of Starobinsky inflation, $H \sim 10^{13} \text{ GeV}$.

in the cosmic time $t = \int a(\eta) d\eta$, and denote the derivatives w.r.t. this time with dots, e.g. $H \equiv \dot{a}/a$.

The accelerated expansion occurs as long as the slow-roll parameters are small. The latter are defined in a usual way,

$$\begin{aligned}\epsilon_1 &\equiv \frac{M_P^2}{2} \left(\frac{V_{,\phi}}{V} \right)^2 = \frac{\phi'^2}{2\mathcal{H}^2 M_P^2}, \\ \epsilon_2 &\equiv \frac{M_P^2 V_{,\phi\phi}}{V}.\end{aligned}\tag{3.3}$$

In the formal limit $V \rightarrow \text{const}$ (equivalently, $\phi' \rightarrow 0$) slow-roll parameters vanish and the background reduces to exact de Sitter space characterized by the scale factor

$$a(\eta) = -1/(H\eta),\tag{3.4}$$

with $H = M_P^{-1} \sqrt{V/3} = \text{const}$.

Now we focus on linear scalar perturbations in the conformal Newtonian gauge,

$$\begin{aligned}ds^2 &= a(\eta)^2 \left[-(1 + 2\Psi(\eta, \mathbf{x})) d\eta^2 + (1 + 2\Phi(\eta, \mathbf{x})) \delta_{ij} dx^i dx^j \right], \\ \phi &= \phi(\eta) + \varphi(\eta, \mathbf{x}).\end{aligned}\tag{3.5}$$

Expanding the action (2.4) to quadratic order one finds,⁶

$$S_{\text{Weyl}}^{(2)} = -\frac{M_P^2}{3m^2} \int d^4x [\Delta(\Psi - \Phi)]^2,\tag{3.6a}$$

$$S_{GR}^{(2)} = \frac{M_P^2}{2} \int d^4x a^2 \left[-6\Phi'^2 + 12\mathcal{H}\Psi\Phi' - 4\Phi\Delta\Psi - 2\Phi\Delta\Phi - 2(\mathcal{H}' + 2\mathcal{H}^2)\Psi^2 \right],\tag{3.6b}$$

$$S_{\phi}^{(2)} = \frac{1}{2} \int d^4x a^2 \left(\varphi'^2 + \varphi\Delta\varphi - a^2 V_{,\phi\phi} \varphi^2 + 6(\phi'\varphi' - V'\varphi)\Phi - 2(\phi'\varphi' + a^2 V_{,\phi} \varphi)\Psi \right),\tag{3.6c}$$

where $\Delta \equiv \partial^i \partial_i$ stands for a spatial Lapacian.

In our analysis we adopt the following strategy. We first consider the limit of exact de Sitter space suggested by the smallness of the slow-roll parameters. In this limit the inflaton decouples from the metric, so the inflaton and the ghost behave as free fields in an external background. At the next step we allow for the finite slow roll-parameters and compute corrections due to the inflaton-metric mixing.

3.1 Exact de Sitter background

We start the analysis with the limit of exact de Sitter space, in which the inflaton decouples from metric perturbations. The latter get excited too since they contain the ghost degree of freedom. We will focus on its scalar polarization, which will be referred to as just “ghost” whenever there is no possible confusion. The quadratic action for the inflaton fluctuations (3.6c) then takes the usual form:

$$S_{\phi}^{(2)} = \frac{1}{2} \int d\eta d^3x a^2 \left[(\varphi')^2 - (\partial_i \varphi)^2 \right].\tag{3.7}$$

⁶Here we omit the terms that vanish due to the background equations (3.2).

One quantizes the inflaton as usual,

$$\varphi(\mathbf{x}, \eta) = \int \frac{d^3k}{(2\pi)^3} (\varphi_{\mathbf{k}}(\eta) a_{\mathbf{k}} + \varphi_{\mathbf{k}}^*(\eta) a_{-\mathbf{k}}^+) e^{i\mathbf{k}\mathbf{x}}, \quad (3.8)$$

with the creation-annihilation operators satisfying

$$[a_{\mathbf{k}}, a_{\mathbf{q}}^+] = (2\pi)^3 \delta^{(3)}(\mathbf{k} - \mathbf{q}). \quad (3.9)$$

The mode functions $\varphi_{\mathbf{k}}$ are positive frequency solutions of equation

$$\varphi_{\mathbf{k}}'' + 2\mathcal{H}\varphi_{\mathbf{k}}' + k^2\varphi_{\mathbf{k}} = 0. \quad (3.10)$$

They have the standard form:

$$\varphi_{\mathbf{k}}^{(dS)} \equiv \frac{H}{\sqrt{2k}} |\eta| \left(1 - \frac{i}{k\eta} \right) e^{-ik\eta}. \quad (3.11)$$

Ghost perturbations. Now let us focus on the metric components. The field Ψ entering eqs. (3.6a) and (3.6b) is non-dynamical since it has only one derivative acting on it. Varying the actions (3.6a) and (3.6b) with respect to Ψ , and performing the Fourier decomposition $\Phi, \Psi \propto e^{i\mathbf{k}\mathbf{x}}$, one finds the following constraint equation:

$$-\frac{2}{3m^2 a^2} k^4 (\Psi - \Phi) + 6\mathcal{H}\Phi' + 2k^2\Phi - 2(2\mathcal{H}^2 + \mathcal{H}')\Psi = 0. \quad (3.12)$$

One observes the appearance of a new time-dependent scale $k_{gh}(\eta) \equiv k^2/(a(\eta)m)$ which also depends on Fourier harmonics. We call this scale *ghost horizon* in what follows. In order to understand the role of this scale, consider the dynamics of a Fourier mode with conformal momentum k (see figure 1). The evolution of this mode starts deep inside the Hubble horizon, when the ghost scale is parametrically enhanced compared to k . In other words, the mode k is outside the ghost horizon. In this regime the leftmost term in eq. (3.12) dominates. Since this term originates from the Weyl tensor squared in the action, one may say that the dynamics is governed by the Weyl term. At the Hubble crossing, $\mathcal{H} = k$, the mode k is still well outside the ghost horizon. As the Universe expands, the ghost scale decreases and eventually gets smaller than the Hubble rate; after that the dynamics of metric perturbations is dominated by terms that originate from the Einstein-Hilbert action.⁷ The ghost horizon gets equal to the Hubble rate after ~ 15 e-foldings since the mode has left the Hubble horizon,

$$\ln \frac{a_{gh, \mathcal{H}}}{a_*} = \ln \sqrt{\frac{H}{m}} \simeq 17 \times \ln \left(\left[\frac{H}{10^{13} \text{ GeV}} \right] \left[\frac{10 \text{ MeV}}{m} \right] \right), \quad (3.13)$$

where $a_{gh, \mathcal{H}}$ is the scale factor at which the Hubble rate starts to dominate over the ghost scale, and a_* is the scale factor at the Hubble crossing.

Eventually, the ghost scale gets smaller than the momentum k and the corresponding mode enters the ghost horizon. This happens after ~ 30 e-foldings after the Hubble crossing,

$$\ln \frac{a_{gh, k}}{a_*} = \ln \frac{H}{m} \simeq 34 \times \ln \left(\left[\frac{H}{10^{13} \text{ GeV}} \right] \left[\frac{10 \text{ MeV}}{m} \right] \right), \quad (3.14)$$

where $a_{gh, k}$ is the scale factor at the ghost horizon crossing. Let us discuss the different dynamical regimes separately.

⁷Recall that we are working in the Einstein frame, in which the theory with R^2 in the action is equivalent to Einstein gravity plus a minimally coupled scalar field.

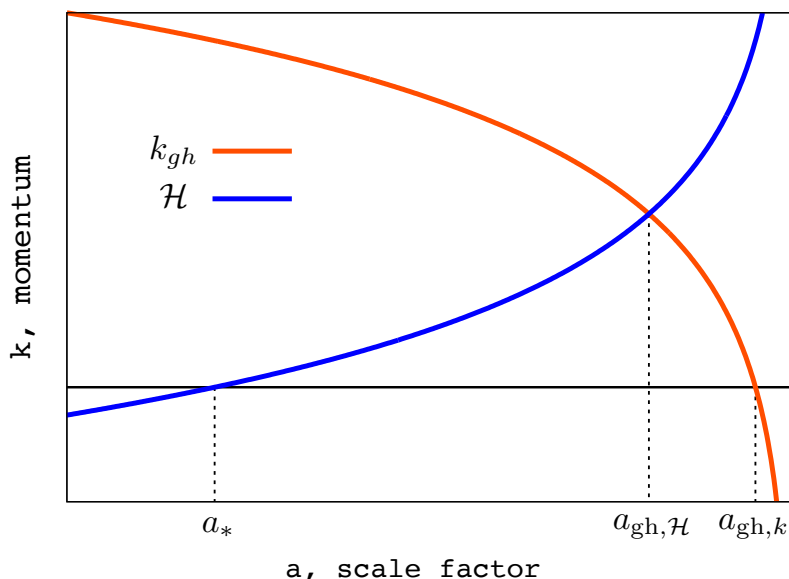


Figure 1. The time dependence of dynamical scales defining the evolution of ghost perturbations: a conformal momentum k (horizontal solid black line), the conformal Hubble parameter \mathcal{H} (solid blue line), and the ghost scale $k_{gh} = k^2/(ma(\eta))$ (solid orange line). For other notations see the main text.

Weyl term domination, $k_{gh} \gg \mathcal{H}, k$. The solution of eq. (3.12) at leading order in $\mathcal{H}/k_{gh}, k/k_{gh}$ is given by

$$\Psi = \Phi, \quad (3.15)$$

which, upon substituting back into eqs. (3.6a) and (3.6b), and canonical normalization $\hat{\Phi} \equiv M_P \sqrt{6} \Phi$, yields the following quadratic action for the Φ field,

$$S_{\hat{\Phi}}^{(2)} \Big|_{k, \mathcal{H} \ll k_{gh}} = \frac{1}{2} \int d^4x a^2 \left[-\hat{\Phi}'^2 + (\partial_i \hat{\Phi})^2 - \frac{4}{3} (2\mathcal{H}^2 + \mathcal{H}') \hat{\Phi}^2 \right]. \quad (3.16)$$

We observe that the scalar part of the massive spin-2 ghost acquires a tachyonic mass on the de Sitter background. In order to quantize the ghost we introduce usual Fourier harmonics analogous to (3.8),

$$\hat{\Phi}(\mathbf{x}, \eta) = \int \frac{d^3k}{(2\pi)^3} (\hat{\Phi}_{\mathbf{k}}(\eta) b_{\mathbf{k}} + \hat{\Phi}_{\mathbf{k}}^*(\eta) b_{-\mathbf{k}}^+) e^{i\mathbf{k}\mathbf{x}} \quad (3.17)$$

with the creation-annihilation operators satisfying⁸

$$[b_{\mathbf{k}}, b_{\mathbf{q}}^+] = -(2\pi)^3 \delta^{(3)}(\mathbf{k} - \mathbf{q}). \quad (3.18)$$

Notice that these operators commute with those introduced in eq. (3.8),

$$[a_{\mathbf{k}}^+, b_{\mathbf{q}}] = [a_{\mathbf{k}}, b_{\mathbf{q}}^+] = 0. \quad (3.19)$$

The mode functions $\hat{\Phi}_{\mathbf{k}}$ satisfy the equation of motion derived from eq. (3.16),

$$\hat{\Phi}_{\mathbf{k}}'' + 2\mathcal{H}\hat{\Phi}_{\mathbf{k}}' + k^2\hat{\Phi}_{\mathbf{k}} - 4\mathcal{H}^2\hat{\Phi}_{\mathbf{k}} = 0, \quad (3.20)$$

⁸We assume that $b_{\mathbf{k}}$ is the annihilation operator for the Bunch-Davies vacuum, i.e. $b_{\mathbf{k}}|0\rangle = 0$.

where we made use $\mathcal{H}' = \mathcal{H}^2$ for de Sitter space. The positive-frequency solution of eq. (3.20) satisfying the Bunch-Davies initial conditions is,

$$\hat{\Phi}_{\mathbf{k}} = \frac{H}{\sqrt{2k}} |\eta| \left(1 - \frac{3i}{k\eta} - \frac{3}{(k\eta)^2} \right) e^{-ik\eta}. \quad (3.21)$$

One observes that the tachyonic mass of the ghost leads to a superhorizon instability developing as the scale factor. This instability signals the breakdown of perturbation theory in the Newtonian gauge. Moreover, the superhorizon growth of ghost perturbations should backfire on the inflaton fluctuations due to the mixing. Nevertheless, let us proceed and see where computations lead to.

Einstein term domination, $\mathcal{H} \gg k_{gh}$. In this case the Weyl term becomes negligible and the solution of eq. (3.12) at next-to-leading order in k_{gh}/\mathcal{H} is given by

$$\Psi = \frac{\Phi'}{\mathcal{H}} + \frac{k^2 \Phi}{3\mathcal{H}^2} - \frac{k_{gh}^2(\eta)}{9\mathcal{H}^2} \left(\frac{\Phi'}{\mathcal{H}} + \frac{k^2 \Phi}{3\mathcal{H}^2} - \Phi \right). \quad (3.22)$$

Switching to real space and plugging the above solution back in the actions (3.6a), (3.6b) yields

$$S_{\Phi}^{(2)} \Big|_{\mathcal{H} \gg k_{gh}} = - \int d^3x d\eta \frac{M_P^2 \Delta^2}{3m^2} \left[\frac{\Phi'}{\mathcal{H}} - \frac{\Delta \Phi}{3\mathcal{H}^2} - \Phi \right]^2. \quad (3.23)$$

The above action implies the following equation of motion in the superhorizon limit, where one can neglect spatial laplacians,

$$\Phi'' + \frac{2}{\eta} \Phi' = 0. \quad (3.24)$$

This equation has a growing solution,

$$\Phi \propto \eta^{-1} \sim a(\eta), \quad (3.25)$$

which accords with the leading superhorizon behavior of the mode function (3.21). Thus, mode functions of Φ do not change as the Hubble rate starts to dominate over the ghost scale. Substituting the solution (3.25) into eq. (3.22) yields, at the leading order in k^2/\mathcal{H}^2 ,

$$\Phi = \Psi. \quad (3.26)$$

Remarkably, the solutions (3.21) and (3.15) do not change when the Weyl terms becomes subdominant. In other words, the ghost perturbations do not “feel” that the Einstein term started to dominate. This is not accidental, the condition (3.15) ((3.26)) reflects the fact that scalar perturbations on exact de Sitter space are conformally flat, in which case the Weyl tensor squared (3.6a) vanishes. The existence of this mode was obtained for the first time in ref. [46]. This mode also appears in some ghost-free non-local generalizations of quadratic gravity, e.g. [47].

To sum up, we have shown that the constraint (3.15) and the solution for ghost mode functions (3.21) are valid all the time during inflation, even though the terms that dominate in the action change.

3.2 Beyond the limit of exact de Sitter space

In this subsection we account for corrections due to the inflaton-ghost mixing. Within our scheme, the metric perturbations obtained in the previous subsection act as sources in the inhomogeneous equation of motion for the inflaton derived from the action (3.6c),

$$\varphi_{\mathbf{k}}'' + 2\mathcal{H}\varphi_{\mathbf{k}}' + k^2\varphi_{\mathbf{k}} = -2\phi'\Phi_{\mathbf{k}}' - 2a^2V_{,\phi}\Phi_{\mathbf{k}}. \quad (3.27)$$

The field Φ is given by eq. (3.21). It has the following leading superhorizon asymptotic:

$$\Phi_{\mathbf{k}} \Big|_{\eta \rightarrow -0} = -\frac{3H^2}{2\sqrt{3} \cdot k k^2 M_P} a(\eta) e^{-ik\eta} \equiv C_{\mathbf{k}} a(\eta), \quad (3.28)$$

where for convenience we switched back to a dimensionless variable and factorized the time dependence. Eq. (3.27) with the source (3.28) has the following solution:

$$\varphi_{\mathbf{k}} = \varphi_{\mathbf{k}}^{(dS)} + \frac{2(\phi'' + \mathcal{H}\phi')}{3\mathcal{H}^2 + \mathcal{H}'} C_{\mathbf{k}} a(\eta). \quad (3.29)$$

Using the conditions $\ddot{\phi} \approx 0$, $\mathcal{H}^2 \approx \mathcal{H}'$ valid in the slow-roll approximation, eq. (3.29) reads

$$\varphi_{\mathbf{k}} = \varphi_{\mathbf{k}}^{(dS)} + \frac{\phi'}{\mathcal{H}} C_{\mathbf{k}} a(\eta). \quad (3.30)$$

One concludes that both the inflaton and ghost perturbations in the Newtonian gauge grow outside the horizon. This growth may be interpreted as a gauge artifact because all perturbations remain bounded in the uniform inflaton gauge.

Uniform inflaton gauge. There are three scalar modes of the metric in the uniform inflaton (comoving slicing) gauge,

$$g_{00} = a^2(1 + 2A), \quad g_{0i} = a^2 \partial_i B, \quad g_{ij} = a^2 \delta_{ij}(1 + 2\mathcal{R}). \quad (3.31)$$

The gauge invariant quantity \mathcal{R} defined as a perturbation of the number of e-foldings on the surfaces of the uniform inflaton field can be retrieved from the modes of the Newtonian gauge via

$$\mathcal{R} = \Phi - \frac{\mathcal{H}}{\phi'} \varphi, \quad (3.32)$$

see appendix A for more detail. One can define mode functions of \mathcal{R} similarly to eq. (3.8) and compute them using the mode functions for φ and Φ (eqs. (3.20) and (3.30)), which yields

$$\mathcal{R}_{\mathbf{k}} = \Phi_{\mathbf{k}} - \frac{\mathcal{H}}{\phi'} \varphi_{\mathbf{k}} = -\frac{\mathcal{H}}{\phi'} \varphi_{\mathbf{k}}^{(dS)} = -\frac{H}{\dot{\phi}} \varphi_{\mathbf{k}}^{(dS)}, \quad (3.33)$$

where $\varphi_{\mathbf{k}}^{(dS)}$ is given by eq. (3.11). One observes that the growing contributions cancel each other in eq. (3.33), so the curvature perturbation coincides with that of the single-field case. Thus, it remains constant outside the horizon analogous to the case of usual gravity [48, 49]. In appendix B we additionally verify that \mathcal{R} is constant outside the horizon by a direct calculation in the uniform inflaton gauge. Note that the constancy of \mathcal{R} in quadratic gravity was proven for the first time in ref. [50] (see also ref. [27]).

Another non-trivial scalar perturbation in the uniform inflaton gauge is the longitudinal part B of the shift vector g_{0i} , see eq. (3.31). The corresponding mode functions can be retrieved from the inflaton perturbation in the Newtonian gauge via

$$B_{\mathbf{k}} = \frac{\varphi_{\mathbf{k}}}{\phi'} = \mathcal{R}_{\mathbf{k}}\eta + \frac{C_{\mathbf{k}}}{H}, \quad (3.34)$$

where in the last equality we have used eqs. (3.30) and (3.33). The first contribution in the r.h.s. of eq. (3.34) (adiabatic mode) vanishes at $\eta \rightarrow 0$. The second contribution is induced by the ghost perturbations and is constant in the limit $\eta \rightarrow 0$,

$$B_{0,\mathbf{k}} \equiv \frac{C_{\mathbf{k}}}{H} = -\frac{\sqrt{3}}{2k^{5/2}} \frac{H}{M_P}. \quad (3.35)$$

The gauge invariant generalisation of B is constructed in appendix A.

Finally, the lapse function A vanishes at superhorizon scales,

$$A = \Psi - \mathcal{H}B - B' = \mathcal{R} - B' = 0, \quad (3.36)$$

where we used eqs. (3.15), (3.33), and (3.34).

We have seen that inflationary perturbations are plagued by superhorizon growth in the conformal Newtonian gauge. This growth, however, should be treated as unphysical since all perturbations are bounded and conserved outside the horizon in the comoving slicing gauge. This suggests two ways to perform calculations:

- 1) remain in the Newtonian gauge and ignore the fact that growing superhorizon modes break perturbation theory. These modes should eventually cancel in gauge-invariant quantities.
- 2) work in the comoving slicing gauge from the beginning.

In this section we followed the first way. In appendix B we perform the analysis directly in the comoving slicing and show that the two approaches yield the same results.

3.3 Observational predictions

The power spectrum of curvature perturbations is given by a standard expression,

$$\mathcal{P}_{\mathcal{R}}(k) = \frac{k^3}{2\pi^2} \lim_{\eta \rightarrow -0} |\mathcal{R}_{\mathbf{k}}(\eta)|^2 = \frac{H^4}{4\pi^2 \dot{\phi}^2} \Big|_{aH=k}, \quad (3.37)$$

where we used eqs. (3.33) and (3.11), and explicitly emphasized that the spectrum above should be evaluated at the horizon crossing.

As for the conserved scalar perturbation B , its power spectrum is given by,

$$\langle |B(\eta, \mathbf{x})|^2 \rangle \xrightarrow{\eta \rightarrow -0} \langle |B_0(\mathbf{x})|^2 \rangle = \int \frac{dk}{k} \frac{\mathcal{P}_B(k)}{k^2}, \quad (3.38)$$

where $\mathcal{P}_B(k)$ is the power spectrum of the shift vector (gradient of B), which closely resembles the power spectrum of tensor modes in single-field inflation,

$$\mathcal{P}_B(k) = \frac{k^5}{2\pi^2} |B_{0,\mathbf{k}}|^2 = \frac{3H^2}{8\pi^2 M_P^2} \Big|_{aH=k}. \quad (3.39)$$

Notice that the amplitude (3.38) is quadratically divergent for long-wavelength perturbations. This is not so worrisome because there exists a natural infrared cutoff set by the Hubble scale. Modes larger than the Hubble horizon cannot be observed locally and can be removed by a suitable gauge transformation, thus they do not contribute to the local energy density. Situation here is similar to the case of primordial gravitational waves, even though their amplitude is only logarithmically divergent for a spectrum close to scale-invariant [51–53].

In order to study the correlation between the fields B and \mathcal{R} we express them through the inflaton φ and the Newtonian potential Φ as in eqs. (3.33) and (3.34). Making use of the decompositions (3.17) and (3.8), and eq. (3.19), we obtain that the correlation between two fields vanishes at superhorizon scales,

$$\langle B(\eta, \mathbf{x}) \mathcal{R}(\eta, \mathbf{x}) \rangle = -\frac{\mathcal{H}}{\phi'^2} \langle |\varphi(\eta, \mathbf{x})|^2 \rangle = \eta \int \frac{dk}{k} \mathcal{P}_{\mathcal{R}} \xrightarrow{\eta \rightarrow -0} 0. \quad (3.40)$$

The power spectrum of tensor modes has been computed in refs. [28–30],

$$\mathcal{P}_h = \frac{2}{\pi^2} \frac{H^2}{M_P^2} \frac{1}{1 + 2\frac{H^2}{m^2}} \Big|_{aH=k} \simeq \frac{1}{\pi^2} \frac{m^2}{M_P^2} \Big|_{aH=k}. \quad (3.41)$$

Note that this result is very similar (up to a numerical factor) to the expression for the spectrum of gravitational waves in Starobinsky inflation [54], which is recovered by changing $m \rightarrow \mu$ (the scalaron mass). Values of the ghost mass lying in the range (2.10) imply an extremely small scalar-to-tensor ratio with no hope for detection,

$$r = \frac{\mathcal{P}_h}{\mathcal{P}_{\mathcal{R}}} = 7 \times 10^{-34} \left[\frac{m}{10 \text{ MeV}} \right]^2 \left[\frac{10^{-10}}{\mathcal{P}_{\mathcal{R}}} \right]. \quad (3.42)$$

The limit $m \rightarrow 0$ corresponds to pure conformal gravity where de Sitter space is equivalent to flat space, in which the amplification of tensor modes does not take place [46]. The spectrum of vector modes was found to be equal to (3.41) [29] and hence is also vanishing.

To sum up, the model predicts the same scalar curvature perturbations as usual inflation, the existence of scalar ghost perturbations and vanishing tensor and vector modes. A non-trivial step in analyzing the scalar sector is the choice of the gauge. We have shown that the conformal Newtonian gauge is unstable because of the tachyonic mass of the ghost. The fluctuations behave well in the comoving slicing, although the lagrangian for perturbations is (approximately) diagonal and can be conventionally quantized only in the Newtonian gauge (see appendix B for more details).

4 After inflation: the cosmological ghost problem

Reheating after inflation takes place during the oscillations of the inflaton field, which cause the production of light particles and can be seen as the inflaton decay. In our setup the ghost is too light to decay and thus reheating in our model is completely identical that of single-field inflation. From now on we denote the Hubble rate by H and use H_{inf} for the Hubble parameter during inflation.

As the Universe expands, the Hubble parameter decreases and eventually gets equal to the ghost mass. The corresponding temperature can be readily estimated,

$$H = 1.66\sqrt{g_*}\frac{T_m^2}{M_P} = m, \Rightarrow$$

$$T_m = \sqrt{\frac{mM_P}{1.66\sqrt{g_*}}} = 3.8 \left[\frac{m}{10^{-1} \text{ eV}}\right]^{1/2} \left[\frac{100}{g_*}\right]^{1/4} \text{ TeV}, \quad (4.1)$$

where g_* denotes the effective number of relativistic degrees of freedom.⁹

One can show that the curvature perturbation \mathcal{R} remains conserved outside the horizon after inflation and sources the adiabatic mode similar to that of standard cosmology (see appendix C for more detail). As for the non-adiabatic mode of B , it remains constant at superhorizon scales as long as $m < H$. In the regime $m > H$ this mode starts to oscillate. Since the ghost cannot decay, the energy of oscillations cannot be spent to produce particles and thus the energy density of the ghost eventually starts to dominate in the Universe. Let us roughly estimate when this happens. A more rigorous derivation is given in appendix C.

As we have seen in the previous section, the canonically normalized ghost field is given, up to constant, by the Newtonian gravitational potential Φ . This field is sourced by the both fields B and \mathcal{R} via

$$\Phi = \mathcal{R} + \mathcal{H}B, \quad (4.2)$$

see appendix A for more detail. Thus, the mode Φ has two contributions: the adiabatic mode related to \mathcal{R}_0 , and the non-adiabatic mode induced by the conserved ghost perturbation B_0 . Let us focus on this non-adiabatic contribution. At superhorizon scales, before the oscillation started, it is given by

$$\Phi_{\text{gh}} = \mathcal{H}B_0. \quad (4.3)$$

As one can see, this mode evolves already at superhorizon scales. The energy density of the ghost is equal to¹⁰

$$\rho_{\text{gh}} = m^2 \langle \hat{\Phi}_{\text{gh}}^2(t, \mathbf{x}) \rangle, \quad (4.4)$$

where $\hat{\Phi}_{\text{gh}} = \sqrt{6}M_P\Phi_{\text{gh}}$ stands for the canonically normalized ghost field. Right before oscillations its value is given by

$$\Phi_{\text{gh}}|_{m=H} = a(t_m)H(t_m)B_0, \quad (4.5)$$

where t_m denotes the transition time at which $m = H(t_m)$. After the Hubble scale has dropped below the ghost mass, the ghost field has the usual solution $\Phi \simeq \text{const} \times \cos(mt + \alpha_0)/a^{3/2}(t)$ (see, e.g. [53]), where α_0 is some phase. Matching this solution to (4.5) and using $H(t_m) = m$ we obtain,

$$\Phi_{\text{gh}} = a(t_m)mB_0 \left(\frac{a(t_m)}{a(t)}\right)^{3/2} \cos(mt + \alpha_0). \quad (4.6)$$

Plugging (4.6) into eq. (4.4) yields,

$$\rho_{\text{gh}} = 6M_P^2m^4a^2(t_m)\langle B_0^2 \rangle \left(\frac{a(t_m)}{a(t)}\right)^3 \cos^2(mt + \alpha_0),$$

$$= \frac{6M_P^2m^4a^5(t_m)\cos^2(mt + \alpha_0)}{a(t)^3} \int \frac{dk}{k} \frac{\mathcal{P}_B}{k^2}, \quad (4.7)$$

⁹Recall that for the Standard Model $g_* = 106.75$ for $T > 200$ GeV.

¹⁰Recall that we sacrificed unitarity for the positive energy density of the ghost.

where we used eq. (3.38) and the integral above should be performed over comoving momenta. This integral is saturated with the infrared cutoff defined by the Hubble scale,

$$k_{\min} = \mathcal{H}(\eta) = aH. \quad (4.8)$$

Neglecting the tilt of the primordial spectrum \mathcal{P}_B , we obtain the local energy density of ghosts inside the horizon,

$$\begin{aligned} \rho_{\text{gh}} &\simeq \frac{3M_P^2 m^4 a^5(t_m)}{2a(t)^3} \frac{1}{a(t)^2 H(t)^2} \mathcal{P}_B \\ &= \frac{3M_P^2 m^4}{2H^2(t)} \left(\frac{T(t)}{T_m} \right)^5 \mathcal{P}_B, \end{aligned} \quad (4.9)$$

where we also averaged over fast oscillations and switched from the scale factor to the temperature T in the last equality.

To find the temperature T_{gh} at which ghosts “overclose” the Universe we equate the energy density (4.9) to the energy density of the background $3H^2 M_P^2$, and use the standard expression for the Hubble rate at radiation domination (see eq. (4.1)). This yields

$$T_{\text{gh}} = [\mathcal{P}_B/2]^{1/3} (M_P^* m)^{1/2} \simeq 260 \left[\frac{H_{\text{inf}}}{10^{13} \text{ GeV}} \right]^{2/3} \left[\frac{m}{10^{-1} \text{ eV}} \right]^{1/2} \left[\frac{100}{g_*} \right]^{1/4} \text{ MeV}. \quad (4.10)$$

Thus, the ghosts start to dominate in the Universe before primordial nucleosynthesis. Assuming that the ghosts are effectively stable, ghost domination can be avoided only for the masses much below 10^{-1} eV, which are excluded by laboratory constraints discussed in section 2. Thus, we conclude that the presence of a stable ghost is incompatible with viable late-time cosmology.¹¹

This result is different from the standard cosmological moduli problem. For a light free scalar field this problem is much less serious, and can be avoided for reasonable values of ghost masses and inflationary Hubble rates. Indeed, let us consider a free field s with a mass m_s much smaller than the Hubble parameter at inflation. Fluctuations of this field $\delta s \sim H_{\text{inf}}$ are generated at inflation and are conserved outside the horizon; the corresponding power spectrum is given by

$$\mathcal{P}_s = \frac{H_{\text{inf}}^2}{(2\pi)^2}. \quad (4.11)$$

In this case the same computation as above would give the following current density fraction:

$$\Omega_s = \frac{\rho_{s,0}}{3H_0^2 M_P^2} \simeq \frac{1}{3M_P^2} \left[\frac{m_s}{H_0} \right]^{1/2} \Omega_\gamma^{3/4} \int_{aH} \frac{dk}{k} \mathcal{P}_s \simeq 5 \times 10^{-3} \left[\frac{H_{\text{inf}}}{10^{11} \text{ GeV}} \right]^2 \left[\frac{m_s}{0.1 \text{ eV}} \right]^{1/2} \left[\frac{N_e}{60} \right], \quad (4.12)$$

where H_0 denotes the current value of the Hubble parameter, Ω_γ stands for the current energy fraction of radiation, $\rho_{s,0}$ stands for the current energy density of the light scalar, and N_e stands the number of e-foldings (we assumed the spectrum \mathcal{P}_s to be scale invariant).

In our case the situation is quite different. The canonically normalized ghost field is not conserved outside the horizon before oscillations. The conserved perturbation is a scalar part of the shift vector in the comoving slicing, which has a different (time-dependent) normalization as compared to a light scalar.

¹¹It should be pointed out that the ghost domination can be avoided in low-scale inflationary scenarios. In appendix C we show that ghosts never dominate in the Universe if the inflationary Hubble rate $H_{\text{inf}} \lesssim 1$ MeV. This requires a very unnatural potential for the inflaton (to say nothing of possible issues with reheating), which is why we do not consider this option here.

5 Discussion and conclusions

In this paper we studied the creation and evolution of cosmological perturbations in quadratic gravity. It has been known for decades that, depending on the quantization prescription, this theory either suffers from the Ostrogradsky instability or is non-unitary. In the first case the theory is pathological both at quantum and classical level, while in the second case it might still be acceptable from the observational point of view if non-unitary processes are suppressed.

As it is, non-unitary quadratic gravity cannot be conventionally interpreted and thus cannot be considered the last word unless the interpretation is found. However, predictions of the theory will not depend on a particular interpretation if the ghost is effectively stable and non-unitary processes have negligible rates during the lifetime of the Universe. We argued that this is the case for ghost masses smaller than 10 MeV. On the other hand, the lower bound is given by laboratory tests of gravity at short scales, which require the ghost to be heavier than 0.1 eV. At face value, this yields a rather wide range of ghost masses which is phenomenologically viable.

We showed that inflation generates the usual curvature perturbations (which are indistinguishable from that of single-field inflation) along with scalar perturbations of the ghost field, which are identified with the longitudinal part of the shift vector in the comoving slicing gauge. The vector and tensor perturbations are suppressed. The scalar ghost fluctuations conserve outside the horizon and their energy density becomes dominant in the Universe shortly after the ghost field started to oscillate. Thus, our study shows that the presence of a light ghost inevitably leads to a ghost dominating phase incompatible with the late-time evolution of our Universe. This result is independent of any interpretation of non-unitary theories.

The ghost domination can be avoided either by allowing the ghost to decay or by assuming that it is not excited during inflation. In the first case the viability of the model depends crucially on the interpretation of processes with negative probabilities. As for the second option, in the case of simplest inflationary scenarios it implies that the ghost mass must be bigger than 10^{13} GeV, which violates the bound $m \lesssim 10^{10}$ GeV required for the naturalness of the electroweak scale in agravity [18]. Our analysis suggests that one would need a rather exotic inflationary scenario in order to have the ghost gapped at inflation and keep the small electroweak scale natural in the adimensional gravity model.

Although we have worked in the framework of renormalizable quadratic gravity, the problem of ghost domination seems to be generic to any theory with the Weyl ghost. The ghost-dominating phase appears to be quite independent of a particular gravity model, and thus should take place in any modified gravity containing the Weyl term with the coupling that corresponds to a stable ghost.

Acknowledgments

We are grateful to D. Levkov, D. Pirtskhalava, V. Rubakov, J. Rubio, and S. Sibiryakov for useful discussions and encouragement. We are indebted to F. Bezrukov, D. Gorbunov, A. Salvio, A. Starobinsky, and A. Strumia for valuable comments on the draft. A. T. acknowledges the support by the Russian Science Foundation grant 14-12-01430 for the analysis of cosmological perturbations at late times. The work of M. I. was supported by the Swiss National Science Foundation. M. I. also acknowledges a partial support by the RFBR grant 14-02-00894.

A Gauge transformations

In this section we outline the relations between perturbations in different gauges. For more detail the reader is advised to consult refs. [55, 56]. Let us consider a generic metric of the form,

$$ds^2 = a^2(\tau) \left\{ -(1 + 2\mathcal{A})d\eta^2 + 2\partial_i\mathcal{B}dx^i d\eta + \left[(1 + 2\mathcal{F})\delta_{ij} + 2 \left(\partial_i\partial_j - \frac{\delta_{ij}\Delta}{3} \right) \mathcal{E} \right] dx^i dx^j \right\}. \quad (\text{A.1})$$

Then the gauge transformations

$$\tilde{x}^\mu = x^\mu + \xi^\mu \quad (\text{A.2})$$

generate the following transformations of the metric components and the inflaton,

$$\begin{aligned} \tilde{\mathcal{A}} &= \mathcal{A} - (\xi^0)' - \mathcal{H}\xi^0, \\ \tilde{\mathcal{F}} &= \mathcal{F} + \frac{1}{3}\Delta\xi - \mathcal{H}\xi^0, \\ \tilde{\mathcal{B}} &= \mathcal{B} + \xi' + \xi^0, \\ \tilde{\mathcal{E}} &= \mathcal{E} + \xi, \\ \tilde{\varphi} &= \varphi - \phi'\xi^0, \end{aligned} \quad (\text{A.3})$$

where $\xi \equiv -\partial_i\xi^i/\Delta$. For a generic matter field the perturbations of its density and velocity potential transform as

$$\begin{aligned} \delta\tilde{\rho}_{\text{tot}} &= \delta\rho_{\text{tot}} - \rho'_{\text{tot}}\xi^0, \\ [(p + \rho)\tilde{v}]_{\text{tot}} &= [(p + \rho)v]_{\text{tot}} + (p + \rho)_{\text{tot}}\xi'. \end{aligned} \quad (\text{A.4})$$

The relations between perturbations in the conformal Newtonian (φ, Φ, Ψ , see eq. (3.5)) and the uniform inflaton (A, B, \mathcal{R} , see eq. (B.1)) gauges are given by

$$\begin{aligned} \Phi &= \mathcal{R} + \mathcal{H}B, \\ \Psi &= A + B' + \mathcal{H}B, \\ \varphi &= \phi'B. \end{aligned} \quad (\text{A.5})$$

Using the transformation rules (A.3) one can define the perturbation B in a gauge invariant way,

$$B^{(\text{g.i.})} = \begin{cases} \mathcal{B} - \mathcal{E}' + \delta\rho_{\text{tot}}/\rho'_{\text{tot}} & \text{in general,} \\ \mathcal{B} - \mathcal{E}' + \varphi/\phi' & \text{at inflation.} \end{cases} \quad (\text{A.6})$$

B Inflationary perturbations in the uniform inflaton gauge

In the main text we have shown that the conformal Newtonian gauge is unstable during inflation since the perturbations undergo an uncontrolled growth outside the horizon. We argued that perturbations behave well in the uniform inflaton gauge, and extracted the corresponding mode functions from those obtained in the Newtonian gauge by means of the gauge transformation (A.5). We argued that this procedure is consistent even though the Newtonian gauge is unstable. In this section we verify this point by a direct calculation in the uniform inflaton gauge.

In the uniform inflaton (comoving slicing) gauge the scalar perturbations take the form

$$\phi = \phi(\eta), \quad g_{00} = -a^2(1 + 2A), \quad g_{0i} = a^2 \partial_i B, \quad g_{ij} = a^2 \delta_{ij}(1 + 2\mathcal{R}). \quad (\text{B.1})$$

The total quadratic action (2.4) then reads,

$$S^{(2)} = M_P^2 \int d^3x d\eta a^2 \left[-\mathcal{R}\Delta\mathcal{R} - 3\mathcal{R}'^2 + 6\mathcal{H}\mathcal{R}'A - 2A\Delta\mathcal{R} - 2\Delta B(\mathcal{H}A - \mathcal{R}') \right. \\ \left. - (\mathcal{H}' + 2\mathcal{H}^2)A^2 - \frac{1}{3a^2m^2}\Delta^2(A - \mathcal{R} + B')^2 \right]. \quad (\text{B.2})$$

The variation of this action with respect to A produces a constraint

$$6\mathcal{H}\mathcal{R}' - 2\Delta\mathcal{R} - 2\Delta B\mathcal{H} - 2A(\mathcal{H}' + 2\mathcal{H}^2) - 2\frac{\Delta^2}{3a^2m^2}(A - \mathcal{R} + B') = 0. \quad (\text{B.3})$$

As in the Newtonian gauge, we observe that the above equation can be solved in two different limits: the ‘‘Weyl term domination’’, characterized by inequality

$$k_{gh} \equiv k^2/(a(\eta)m) \gg k, \mathcal{H},$$

and the ‘‘Einstein term domination’’, at which the above inequality changes to the opposite one. Solving the constraint (B.3) in the regime of the Weyl term domination ($k_{gh} \gg k, \mathcal{H}$) and plugging the solution back yields the following quadratic action:

$$S^{(2)} \Big|_{k_{gh} \gg \mathcal{H}, k} = M_P^2 \int d^3x d\eta a^2 \left[-3\mathcal{R}'^2 - 3\mathcal{R}\Delta\mathcal{R} - 4(\mathcal{H}' + 2\mathcal{H}^2)\mathcal{R}^2 - (\mathcal{H}' + 2\mathcal{H}^2)(B'^2 + B\Delta B) \right. \\ \left. - 6\mathcal{H}\mathcal{R}'B' - 6\mathcal{H}\mathcal{R}\Delta B + 2(\mathcal{H}' + 2\mathcal{H}^2)\mathcal{R}B' \right]. \quad (\text{B.4})$$

In order to quantize the fields \mathcal{R} and B we have to diagonalize the above action inside the Hubble horizon. This amounts to introducing new fields $\tilde{\chi}$ and $\tilde{\Phi}$ via

$$\tilde{\chi} = \phi' B, \quad (\text{B.5a})$$

$$\tilde{\Phi} = \sqrt{6}M_P(\mathcal{R} + \mathcal{H}B), \quad (\text{B.5b})$$

for which the action (B.4) takes the following form inside the horizon at the leading order in the slow roll parameters,

$$S^{(2)} \Big|_{k_{gh} \gg k \gg \mathcal{H}} = \frac{1}{2} \int d^3x d\eta a^2 \left[(\tilde{\chi}'^2 - (\partial_i \tilde{\chi})^2) - (\tilde{\Phi}'^2 - (\partial_i \tilde{\Phi})^2) \right], \quad (\text{B.6})$$

Looking at eq. (A.5) we see that eq. (B.5) is nothing but a gauge transformation, which brings us back to the Newtonian gauge; the field $\tilde{\chi}$ coincides with the inflaton, the field $\tilde{\Phi}$ — with the canonically normalized Newtonian potential. These fields represent two independent (inside the horizon) degrees of freedom. This implies that quantization can conventionally be performed only in the Newtonian gauge.

Now we check that all perturbations in the uniform inflaton gauge get frozen outside the horizon. The action (B.4) implies the following equations of motion,

$$\mathcal{R}'' + 2\mathcal{H}\mathcal{R}' - \Delta\mathcal{R} + \frac{4}{3}(\mathcal{H}' + 2\mathcal{H}^2)(B' - \mathcal{R}) + \mathcal{H}(B'' - \Delta B) = 0, \quad (\text{B.7}) \\ (\mathcal{H}' + 2\mathcal{H}^2)(B'' - \Delta B) + (\mathcal{H}'' + 6\mathcal{H}'\mathcal{H} + 4\mathcal{H}^3)(B' - \mathcal{R}) + 2(\mathcal{H}' + 2\mathcal{H}^2)\mathcal{R}' + 3\mathcal{H}(\mathcal{R}'' - \Delta\mathcal{R}) = 0.$$

Naively the above equations are regular in the limit of exact de Sitter space. In this case the two field equations are equivalent, which is a consequence of the fact that \mathcal{R} becomes a gauge mode, and thus we are left only with only one ghost degree of freedom.

Now let us find a non-trivial solution to the equations eqs. (B.7) in the slow-roll limit. Upon summing and subtracting these equations from each other they can be brought to the following form,

$$\begin{aligned} (\mathcal{H}' - \mathcal{H}^2) \left[B'' - \Delta B + 2\mathcal{R}' + (B' - \mathcal{R}) \frac{\mathcal{H}'' + 2\mathcal{H}\mathcal{H}' - 4\mathcal{H}^3}{\mathcal{H}' - \mathcal{H}^2} \right] &= 0, \\ \frac{(\mathcal{H}' - \mathcal{H}^2)}{\mathcal{H}' + 2\mathcal{H}^2} \left[\mathcal{R}'' - \Delta\mathcal{R} - (B' - \mathcal{R}) \frac{3\mathcal{H}\mathcal{H}'' - 4\mathcal{H}'^2 + 2\mathcal{H}^2\mathcal{H}' - 4\mathcal{H}^4}{3(\mathcal{H}' - \mathcal{H}^2)} \right] &= 0. \end{aligned} \quad (\text{B.8})$$

Note that no slow-roll approximation has been made so far. Now assume that the modes are outside the Hubble horizon, hence the terms with spacial laplacians can be neglected. Next we combine the above equations in order to obtain a single homogeneous equation on \mathcal{R} , and assume a quasi-de Sitter background with $a \propto \eta^{1/(\epsilon_1-1)}$. One obtains,

$$\mathcal{R}''' + \mathcal{R}'' \frac{4}{(\epsilon_1 - 1)\eta} + \frac{2(3 - \epsilon_1)}{(\epsilon_1 - 1)^2 \eta^2} \mathcal{R}' = 0. \quad (\text{B.9})$$

The solution at the leading order in ϵ_1 is given by a constant mode and two modes that vanish in the limit $\eta \rightarrow 0$,

$$\mathcal{R} = \{ \text{const}, \quad \eta^{4+2\epsilon_1}, \quad \eta^{3+2\epsilon_1} \}. \quad (\text{B.10})$$

The constant mode $\mathcal{R} = \mathcal{R}_0$ implies the following solution for the mode B ,

$$\begin{aligned} \mathcal{R} = B' = \mathcal{R}_0 = \text{const}, \\ \Rightarrow B = \mathcal{R}_0 \eta + B_0, \end{aligned} \quad (\text{B.11})$$

where B_0 is a constant. The decaying modes of \mathcal{R} (B.10) source only decaying solutions for B .

In the limit of the Einstein term domination ($k_{gh} \ll \mathcal{H}$) one finds the following action at first order in k^2/\mathcal{H}^2 , k_{gh}^2/\mathcal{H}^2 , ϵ_1 (see eq. (3.3)),

$$S^{(2)} \Big|_{k_{gh} \ll \mathcal{H}} = M_P^2 \int d^3x d\eta a^2 \left[\epsilon_1 \mathcal{R}'^2 - \frac{\Delta^2}{3m^2 a^2} \left(B' + \frac{3\mathcal{H}\mathcal{R}'}{\mathcal{H}' + 2\mathcal{H}^2} - \mathcal{R} \right)^2 \right]. \quad (\text{B.12})$$

It is trivial to show that the solution (B.11) holds in this case too.

To sum up, we have shown that the solution (B.11) coincides with the one obtained by gauge transforming the solutions from the Newtonian gauge (see eqs. (3.33) and (3.34)).

C Cosmological perturbations at late times

C.1 Main equations: qualitative analysis

In this section we consider the dynamics of linear cosmological perturbations after inflation and reheating. Previous results suggest us to work from the beginning in the comoving slicing gauge where we expect perturbations to be free of unphysical instabilities. For simplicity, we will focus on the case of a single component matter, which is known to be a very good approximation since it describes perturbations of a component that dominates at a given

cosmological epoch. This approximation will be enough for our purposes. In practice, we use the energy-momentum tensor

$$T_{\mu\nu} = (\rho + p)v_\mu v_\nu + pg_{\mu\nu}, \quad (\text{C.1})$$

assuming a constant equation of state $p = c_s^2 \rho$. Since the Weyl term vanishes on the FRW background, assuming the metric (3.2) and the ansatz $\rho = \bar{\rho}(\eta)$, $p = \bar{p}(\eta)$ we obtain the usual set of the Friedman and continuity equations,

$$\begin{aligned} 3M_P^2 \mathcal{H}^2 &= a^2 \bar{\rho}, \\ \bar{\rho}' + 3\mathcal{H}(1 + c_s^2)\bar{\rho} &= 0. \end{aligned} \quad (\text{C.2})$$

At the level of perturbations, we assume the metric on the comoving slicing (B.1) and the following perturbations of the energy-momentum tensor,

$$\begin{aligned} \rho &= \bar{\rho} + \delta\rho, \\ p &= \bar{p} + \delta p, \\ v_\mu &= (-a(1 + \Psi), 0), \\ v^\mu &= \frac{1}{a}((1 - \Psi), -\partial_i B), \end{aligned} \quad (\text{C.3})$$

where we have used the definition of the comoving slicing gauge $\delta T_i^0 = 0$ [55]. Substituting the metric (3.31) and the energy momentum tensor components (C.3) into (2.6) we find the following set of modified Einstein equations,

$$-\Delta(\mathcal{R} + \mathcal{H}B) + 3\mathcal{H}\mathcal{R}' - 3\mathcal{H}^2 A - \frac{\Delta^2}{3a^2 m^2}(A + B' - \mathcal{R}) = \frac{a^2}{2M_P} \delta\rho, \quad (\text{C.4a})$$

$$-\mathcal{R}' + \mathcal{H}A - \frac{\Delta}{3a^2 m^2}(A' + B'' - \mathcal{R}') = 0, \quad (\text{C.4b})$$

$$\begin{aligned} -\mathcal{R}'' - 2\mathcal{H}\mathcal{R}' + \mathcal{H}A' + \frac{1}{3}\Delta(\mathcal{R} + B' + A + 2\mathcal{H}B) \\ + (\mathcal{H}^2 + 2\mathcal{H}')A - \frac{\Delta^2}{9m^2 a^2}(A + B' - \mathcal{R}) - \frac{a^2}{2M_P^2} \delta p = 0, \end{aligned} \quad (\text{C.4c})$$

$$A + B' + 2\mathcal{H}B + \mathcal{R} + \frac{1}{a^2 m^2} \left(\partial_\eta^2 - \frac{\Delta}{3} \right) (A + B' - \mathcal{R}) = 0. \quad (\text{C.4d})$$

The conservation of the energy momentum tensor yields two more equations,

$$A = -\frac{\delta p}{(\bar{p} + \bar{\rho})}, \quad (\text{C.5a})$$

$$A' = c_s^2 \Delta B + 3c_s^2 \mathcal{R}', \quad (\text{C.5b})$$

where in order to obtain the last equation we used $\delta p = c_s^2 \delta\rho$.

Upon inspection of eqs. (C.4) and (C.5) one notices that there are three equations that form a closed system for the variables B , A and \mathcal{R} ,

$$\mathcal{R}' - \mathcal{H}A = -\frac{\Delta}{3a^2 m^2}(B'' + A' - \mathcal{R}'), \quad (\text{C.6a})$$

$$A + B' + 2\mathcal{H}B + \mathcal{R} + \frac{1}{a^2 m^2} \left(\partial_\eta^2 - \frac{\Delta}{3} \right) (A + B' - \mathcal{R}) = 0, \quad (\text{C.6b})$$

$$\begin{cases} A' = c_s^2(\Delta B - 3\mathcal{R}') & \text{if } c_s^2 \neq 0, \\ A = 0 & \text{if } c_s^2 = 0. \end{cases} \quad (\text{C.6c})$$

Now let us study eqs. (C.6) in the superhorizon limit where we can neglect spatial laplacians. We also assume the scale factor $a(\eta) \propto \eta^p$. This yields in the following system:

$$\mathcal{R}' - \frac{p}{\eta}A = 0, \tag{C.7a}$$

$$A + B' + 2\frac{p}{\eta}B + \mathcal{R} + \frac{1}{a^2m^2}(A'' + B''' - \mathcal{R}'') = 0, \tag{C.7b}$$

$$\begin{cases} A' = -3c_s^2\mathcal{R}' & \text{if } c_s^2 \neq 0, \\ A = 0 & \text{if } c_s^2 = 0. \end{cases} \tag{C.7c}$$

The mode $\mathcal{R} = \mathcal{R}_0 = \text{const}$ induces the standard adiabatic mode,

$$\begin{aligned} A &= 0, \\ B &= -\frac{\mathcal{R}_0\eta}{1+2p}. \end{aligned} \tag{C.8}$$

Remarkably, this solution is valid regardless the value of the ghost mass, i.e. for both regimes $H \gg m$ and $H \ll m$.

In order to study the dynamics of the isocurvature mode of B we consider eq. (C.7b) with a vanishing source $\mathcal{R} = 0$ (which also implies $A = 0$). We also assume, for definiteness, the radiation domination stage with $p = 1$. In this case eq. (C.7b) reduces to

$$B''' + a^2m^2 \left(B' + \frac{2}{\eta}B \right) = 0. \tag{C.9}$$

In the limit $m \rightarrow 0$ the above equations has three solutions, $B = \text{const}$, $B \propto \eta$ and $B \propto \eta^2$, but the last two solutions are negligible in the limit $\eta \rightarrow 0$. This implies that the inflationary mode $B = B_0$ is conserved.

Once the Hubble rate drops below the ghost mass, the field B starts to oscillate. Eq. (C.9) can be rewritten as

$$\left[\Phi_{\text{gh}}'' + \frac{2}{\eta}\Phi_{\text{gh}}' + a^2m^2\Phi_{\text{gh}} \right]' = 0, \quad \text{where } \Phi_{\text{gh}} = \mathcal{H}B = \frac{B}{\eta}, \tag{C.10}$$

so it takes the form of a usual Klein-Gordon-Fock equation for the massive field at superhorizon scales. Eq. (C.10) describes the behavior of a contribution to the Newtonian potential that is induced by a non-adiabatic mode of B . The initial conditions for an oscillating solution of eq. (C.10) are defined by a constant B_0 .

C.2 Ghost energy density

As we have seen in the previous section, the field B has a constant non-adiabatic mode at superhorizon scales as long as $m < H$. Once the Hubble rate gets smaller than the ghost mass, the field B starts to oscillate. Since the interaction of the ghost with matter is suppressed, the production of particles does not occur and the oscillation energy of the ghost field starts to dominate the cosmic expansion. Neglecting the coupling to matter and using eqs. (3.6a) and (3.6b) it can be shown that at the leading order in the Hubble rate the quadratic action of the ghost takes the following form,¹²

$$S = -3M_P^2 \int d^3x dt a^3 \left(\dot{\Phi}_{\text{gh}}^2 - \frac{(\partial_i \Phi_{\text{gh}})^2}{a^2} - m^2 \Phi_{\text{gh}}^2 \right). \tag{C.11}$$

¹²Formally, in deriving the quadratic action (3.6b) we made use of the background equations of motion for the inflaton. In the case of an arbitrary matter content the result will be the same, which can easily be shown e.g. by writing the macroscopic matter action along the lines of ref. [57].

The non-adiabatic constant mode of B defines the following asymptotics of Φ_{gh} in the limit $t \rightarrow 1/m$:

$$\Phi_{\text{gh}} = \mathcal{H}B_0 \rightarrow \frac{a_\gamma}{2t^{1/2}}B_0, \quad (\text{C.12})$$

where we used the scale factor $a(t) = a_\gamma \cdot t^{1/2}$ at radiation domination. A concrete expression for a_γ will not be relevant for us. The action (C.11) yields the following equation of motion:

$$\ddot{\Phi}_{\text{gh}} + 3H\dot{\Phi}_{\text{gh}} + \left(-\frac{\Delta}{a^2} + m^2\right)\Phi_{\text{gh}} = 0. \quad (\text{C.13})$$

For cosmologically relevant modes the mass term always dominates, hence one can safely neglect the spatial laplacian. Then the solution of eq. (C.13) satisfying the initial behavior (C.12) is given by,

$$\Phi_{\text{gh}} = -\frac{\pi a_\gamma B_0 \sqrt[4]{m} Y_{1/4}(mt)}{2\sqrt[4]{2}\Gamma\left(\frac{1}{4}\right) t^{1/4}}, \quad (\text{C.14})$$

where $Y_{1/4}(x)$ is the Bessel function of the second kind, and $\Gamma(x)$ is the Euler gamma function. At $t \gg 1/m$ the solution (C.14) has the following asymptotic:

$$\Phi_{\text{gh}} \approx \frac{\pi a_\gamma B_0 \sqrt[4]{m}}{2\sqrt[4]{2}\Gamma\left(\frac{1}{4}\right) t^{1/4}} \sqrt{\frac{2}{\pi(mt)}} \cos\left(mt + \frac{\pi}{8}\right) = \frac{\sqrt{\pi}}{2^{3/4}\Gamma\left(\frac{1}{4}\right)} B_0 \frac{a_\gamma^{5/2}}{m^{1/4}a(t)^{3/2}} \cos\left(mt + \frac{\pi}{8}\right). \quad (\text{C.15})$$

In order to estimate the energy density of the ghost we rewrite the action (C.11) in comoving coordinates $\tilde{x} \equiv a(t)x$,

$$S = -3M_P^2 \int d^3\tilde{x} dt \left(\dot{\Phi}_{\text{gh}}^2 - (\tilde{\partial}_i \Phi_{\text{gh}})^2 - m^2 \Phi_{\text{gh}}^2 \right), \quad (\text{C.16})$$

and compute the corresponding energy functional using the decomposition (3.17) and the negative-norm prescription (3.18), which yields,

$$E = \int d^3\tilde{x} \rho_{\text{gh}} = \int d^3\tilde{x} \left[3M_P^2 \int \frac{dq}{q} \Phi_{\text{gh}}^2(q) (\omega_q^2 + q^2 + m^2) \right], \quad (\text{C.17})$$

where ω_q and q are the comoving frequency and momentum. Plugging the amplitude (C.15) in eq. (C.17) and using the dispersion relation $\omega_q^2 = m^2 + q^2 \approx m^2$ one finds,

$$\begin{aligned} \rho_{\text{gh}} &= 6m^2 M_P^2 \int \frac{dq}{q} \Phi_{\text{gh}}^2(q), \\ &= \frac{\pi}{2^{3/2}\Gamma^2\left(\frac{1}{4}\right)} \frac{a_\gamma^5}{m^{1/2}a(t)^3} 6m^2 M_P^2 \cos^2\left(mt + \frac{\pi}{8}\right) \times \left[\int \frac{dq}{q} B_0^2(q) \right], \end{aligned} \quad (\text{C.18})$$

By introducing the transition time

$$t_m : H(t_m) = m,$$

Eq. (C.18) can be rewritten as

$$\rho_{\text{gh}} = \frac{\pi}{\Gamma^2\left(\frac{1}{4}\right)} \frac{6m^4 M_P^2 a^5(t_m)}{a(t)^3} \cos^2\left(mt + \frac{\pi}{8}\right) \times \left[\int \frac{dq}{q} B_0^2(q) \right]. \quad (\text{C.19})$$

This expression coincides with our rough estimate (4.7) up to a factor $\pi/\Gamma^2\left(\frac{1}{4}\right) \approx 0.24$.

In section 4 we showed that for reasonable values of ghost masses $m \gtrsim 0.1 \text{ eV}$ and Hubble rates at inflation $H \sim 10^{13} - 10^{16} \text{ GeV}$ the ghost energy density becomes dominant before primordial nucleosynthesis. Let us assume now that the inflation happened at very low energies and estimate for which value of the inflationary Hubble rate the ghost domination does not happen. First relevant observation is that the solution (C.15) holds at matter domination, which implies that the expression (C.19) does not change either. Then, using the expression for the scale factor at radiation domination $a = a_\gamma t^{1/2} = \sqrt[4]{\Omega_\gamma} \sqrt{2H_0 t}$ we find the current energy fraction of the ghosts,

$$\begin{aligned} \Omega_{\text{gh}} &\equiv \frac{\rho_{\text{gh},0}}{3M_P^2 H_0^2} = \frac{\pi}{\Gamma^2\left(\frac{1}{4}\right)} \mathcal{P}_B \Omega_\gamma^{5/4} \left[\frac{m}{H_0}\right]^{3/2} \\ &\simeq 0.1 \left[\frac{H_{\text{inf}}}{4 \text{ MeV}}\right]^2 \left[\frac{m}{10^{-1} \text{ eV}}\right]^{3/2}, \end{aligned} \tag{C.20}$$

where H_{inf} denotes the Hubble parameter at inflation, Ω_γ — the current energy fraction of radiation, H_0 — current Hubble rate. This expression should be contrasted with eq. (4.12). Compared to a light moduli, given the same masses, the ghost field requires a significantly lower scale of inflation in order not to “overclose” the Universe.

References

- [1] J. Polchinski, *String theory*, Cambridge University Press, Cambridge U.K. (1998).
- [2] P. Hořava, *Quantum gravity at a Lifshitz point*, *Phys. Rev. D* **79** (2009) 084008 [[arXiv:0901.3775](#)] [[INSPIRE](#)].
- [3] A.O. Barvinsky, D. Blas, M. Herrero-Valea, S.M. Sibiryakov and C.F. Steinwachs, *Renormalization of Hořava gravity*, *Phys. Rev. D* **93** (2016) 064022 [[arXiv:1512.02250](#)] [[INSPIRE](#)].
- [4] D. Blas, O. Pujolàs and S. Sibiryakov, *On the extra mode and inconsistency of Hořava gravity*, *JHEP* **10** (2009) 029 [[arXiv:0906.3046](#)] [[INSPIRE](#)].
- [5] D. Blas, O. Pujolàs and S. Sibiryakov, *Models of non-relativistic quantum gravity: the good, the bad and the healthy*, *JHEP* **04** (2011) 018 [[arXiv:1007.3503](#)] [[INSPIRE](#)].
- [6] D. Blas, O. Pujolàs and S. Sibiryakov, *Consistent extension of Hořava gravity*, *Phys. Rev. Lett.* **104** (2010) 181302 [[arXiv:0909.3525](#)] [[INSPIRE](#)].
- [7] B. Audren, D. Blas, M.M. Ivanov, J. Lesgourgues and S. Sibiryakov, *Cosmological constraints on deviations from Lorentz invariance in gravity and dark matter*, *JCAP* **03** (2015) 016 [[arXiv:1410.6514](#)] [[INSPIRE](#)].
- [8] K. Yagi, D. Blas, E. Barausse and N. Yunes, *Constraints on Einstein-Ether theory and Hořava gravity from binary pulsar observations*, *Phys. Rev. D* **89** (2014) 084067 [*Erratum ibid.* **D 90** (2014) 069901] [*Erratum ibid.* **D 90** (2014) 069902] [[arXiv:1311.7144](#)] [[INSPIRE](#)].
- [9] M.M. Ivanov and S. Sibiryakov, *UV-extending ghost inflation*, *JCAP* **05** (2014) 045 [[arXiv:1402.4964](#)] [[INSPIRE](#)].
- [10] D. Blas, M.M. Ivanov and S. Sibiryakov, *Testing Lorentz invariance of dark matter*, *JCAP* **10** (2012) 057 [[arXiv:1209.0464](#)] [[INSPIRE](#)].
- [11] D. Blas and S. Sibiryakov, *Technically natural dark energy from Lorentz breaking*, *JCAP* **07** (2011) 026 [[arXiv:1104.3579](#)] [[INSPIRE](#)].
- [12] D. Blas and E. Lim, *Phenomenology of theories of gravity without Lorentz invariance: the preferred frame case*, *Int. J. Mod. Phys. D* **23** (2015) 1443009 [[arXiv:1412.4828](#)] [[INSPIRE](#)].

- [13] K.S. Stelle, *Classical gravity with higher derivatives*, *Gen. Rel. Grav.* **9** (1978) 353 [INSPIRE].
- [14] K.S. Stelle, *Renormalization of higher derivative quantum gravity*, *Phys. Rev. D* **16** (1977) 953 [INSPIRE].
- [15] E.S. Fradkin and A.A. Tseytlin, *Renormalizable asymptotically free quantum theory of gravity*, *Phys. Lett. B* **104** (1981) 377 [INSPIRE].
- [16] E.S. Fradkin and A.A. Tseytlin, *Renormalizable asymptotically free quantum theory of gravity*, *Nucl. Phys. B* **201** (1982) 469 [INSPIRE].
- [17] I.G. Avramidi and A.O. Barvinsky, *Asymptotic freedom in higher derivative quantum gravity*, *Phys. Lett. B* **159** (1985) 269 [INSPIRE].
- [18] A. Salvio and A. Strumia, *Agravity*, *JHEP* **06** (2014) 080 [arXiv:1403.4226] [INSPIRE].
- [19] M.B. Einhorn and D.R.T. Jones, *Naturalness and dimensional transmutation in classically scale-invariant gravity*, *JHEP* **03** (2015) 047 [arXiv:1410.8513] [INSPIRE].
- [20] J.F. Donoghue, *A conformal model of gravitons*, arXiv:1609.03524 [INSPIRE].
- [21] M. Ostrogradski, *Mémoire sur les équations différentielles relatives au problème des isopérimètres* (in French), *Mem. Ac. St. Petersburg* **VI** (1850) 385.
- [22] R.P. Woodard, *Ostrogradsky's theorem on Hamiltonian instability*, *Scholarpedia* **10** (2015) 32243 [arXiv:1506.02210] [INSPIRE].
- [23] A. Pais and G.E. Uhlenbeck, *On field theories with nonlocalized action*, *Phys. Rev.* **79** (1950) 145 [INSPIRE].
- [24] A. Salvio and A. Strumia, *Quantum mechanics of 4-derivative theories*, *Eur. Phys. J. C* **76** (2016) 227 [arXiv:1512.01237] [INSPIRE].
- [25] J. Maldacena, *Einstein gravity from conformal gravity*, arXiv:1105.5632 [INSPIRE].
- [26] R.P. Feynman, *Negative probability*, PRINT-84-0846, California Institute of Technology, U.S.A. (1984) [INSPIRE].
- [27] N. Deruelle, M. Sasaki, Y. Sendouda and A. Youssef, *Inflation with a Weyl term, or ghosts at work*, *JCAP* **03** (2011) 040 [arXiv:1012.5202] [INSPIRE].
- [28] T. Clunan and M. Sasaki, *Tensor ghosts in the inflationary cosmology*, *Class. Quant. Grav.* **27** (2010) 165014 [arXiv:0907.3868] [INSPIRE].
- [29] Y.S. Myung and T. Moon, *Primordial massive gravitational waves from Einstein-Chern-Simons-Weyl gravity*, *JCAP* **08** (2014) 061 [arXiv:1406.4367] [INSPIRE].
- [30] Y.S. Myung and T. Moon, *Scale-invariant tensor spectrum from conformal gravity*, *Mod. Phys. Lett. A* **30** (2015) 1550172 [arXiv:1501.01749] [INSPIRE].
- [31] G.D. Coughlan, W. Fischler, E.W. Kolb, S. Raby and G.G. Ross, *Cosmological problems for the Polonyi potential*, *Phys. Lett. B* **131** (1983) 59 [INSPIRE].
- [32] A.A. Starobinsky, *A new type of isotropic cosmological models without singularity*, *Phys. Lett. B* **91** (1980) 99 [INSPIRE].
- [33] A.A. Starobinsky, *Nonsingular model of the universe with the quantum-gravitational de Sitter stage and its observational consequences*, in *Proc. of the Second Seminar "Quantum Theory of Gravity"*, Moscow Russia October 13–15 1981, INR Press, Moscow Russia (1982), pg. 58, reprinted in *Quantum Gravity*, M.A. Markov and P.C. West eds., Plenum Publ. Co., New York U.S.A. (1984), pg. 103.
- [34] T. Ruzmaikina and A. Ruzmaikin, *Quadratic corrections to the Lagrangian density of the gravitational field and the singularity*, *JETP Lett.* **30** (1970) 372 [*Pisma Zh. Eksp. Teor. Fiz.* **57** (1970) 680].
- [35] J.D. Barrow and S. Hervik, *On the evolution of universes in quadratic theories of gravity*, *Phys. Rev. D* **74** (2006) 124017 [gr-qc/0610013] [INSPIRE].

- [36] M. Ivanov and A.V. Toporensky, *Stable super-inflating cosmological solutions in $f(R)$ -gravity*, *Int. J. Mod. Phys. D* **21** (2012) 1250051 [[arXiv:1112.4194](#)] [[INSPIRE](#)].
- [37] G. Magnano, M. Ferraris and M. Francaviglia, *Nonlinear gravitational Lagrangians*, *Gen. Rel. Grav.* **19** (1987) 465 [[INSPIRE](#)].
- [38] A. De Felice and S. Tsujikawa, *$f(R)$ theories*, *Living Rev. Rel.* **13** (2010) 3 [[arXiv:1002.4928](#)] [[INSPIRE](#)].
- [39] K. Kannike et al., *Dynamically induced Planck scale and inflation*, *JHEP* **05** (2015) 065 [[arXiv:1502.01334](#)] [[INSPIRE](#)].
- [40] B. Grinstein, D. O'Connell and M.B. Wise, *Causality as an emergent macroscopic phenomenon: the Lee-Wick $O(N)$ model*, *Phys. Rev. D* **79** (2009) 105019 [[arXiv:0805.2156](#)] [[INSPIRE](#)].
- [41] D.S. Gorbunov and A.G. Panin, *Scalaron the mighty: producing dark matter and baryon asymmetry at reheating*, *Phys. Lett. B* **700** (2011) 157 [[arXiv:1009.2448](#)] [[INSPIRE](#)].
- [42] A.A. Starobinsky and Ya.B. Zeldovich, *Rate of particle production in gravitational fields*, *JETP Lett.* **26** (1977) 252 [*Pisma Zh. Eksp. Teor. Fiz.* **26** (1977) 373].
- [43] A. Vilenkin, *Classical and quantum cosmology of the Starobinsky inflationary model*, *Phys. Rev. D* **32** (1985) 2511 [[INSPIRE](#)].
- [44] A.A. Geraci, S.J. Smullin, D.M. Weld, J. Chiaverini and A. Kapitulnik, *Improved constraints on non-Newtonian forces at 10 microns*, *Phys. Rev. D* **78** (2008) 022002 [[arXiv:0802.2350](#)] [[INSPIRE](#)].
- [45] W.-H. Tan et al., *New test of the gravitational inverse-square law at the submillimeter range with dual modulation and compensation*, *Phys. Rev. Lett.* **116** (2016) 131101 [[INSPIRE](#)].
- [46] A.A. Starobinsky, *Evolution of small excitation of isotropic cosmological models with one loop quantum gravitation corrections* (in Russian), *Zh. Eksp. Teor. Fiz.* **34** (1981) 460 [[INSPIRE](#)].
- [47] A.S. Koshelev, L. Modesto, L. Rachwal and A.A. Starobinsky, *Occurrence of exact R^2 inflation in non-local UV-complete gravity*, *JHEP* **11** (2016) 067 [[arXiv:1604.03127](#)] [[INSPIRE](#)].
- [48] J.M. Bardeen, P.J. Steinhardt and M.S. Turner, *Spontaneous creation of almost scale-free density perturbations in an inflationary universe*, *Phys. Rev. D* **28** (1983) 679 [[INSPIRE](#)].
- [49] D.S. Salopek and J.R. Bond, *Nonlinear evolution of long wavelength metric fluctuations in inflationary models*, *Phys. Rev. D* **42** (1990) 3936 [[INSPIRE](#)].
- [50] A.A. Starobinsky, S. Tsujikawa and J. Yokoyama, *Cosmological perturbations from multifield inflation in generalized Einstein theories*, *Nucl. Phys. B* **610** (2001) 383 [[astro-ph/0107555](#)] [[INSPIRE](#)].
- [51] A.A. Starobinsky, *Spectrum of relict gravitational radiation and the early state of the universe*, *JETP Lett.* **30** (1979) 682 [*Pisma Zh. Eksp. Teor. Fiz.* **30** (1979) 719] [[INSPIRE](#)].
- [52] B. Allen, *The stochastic gravity wave background in inflationary universe models*, *Phys. Rev. D* **37** (1988) 2078 [[INSPIRE](#)].
- [53] D.S. Gorbunov and V.A. Rubakov, *Introduction to the theory of the early universe: cosmological perturbations and inflationary theory*, World Scientific, Hackensack U.S.A. (2011).
- [54] A.A. Starobinsky, *The perturbation spectrum evolving from a nonsingular initially de-Sitter cosmology and the microwave background anisotropy*, *Sov. Astron. Lett.* **9** (1983) 302 [[INSPIRE](#)].
- [55] W. Hu, *Covariant linear perturbation formalism*, [astro-ph/0402060](#) [[INSPIRE](#)].
- [56] H. Kurki-Suonio, *Cosmological perturbation theory*, <http://www.helsinki.fi/~hkurkisu/CosPer.pdf>, September 30 2012.
- [57] V.F. Mukhanov, H.A. Feldman and R.H. Brandenberger, *Theory of cosmological perturbations. Part 1. Classical perturbations. Part 2. Quantum theory of perturbations. Part 3. Extensions*, *Phys. Rept.* **215** (1992) 203 [[INSPIRE](#)].